

## Combined BIMA and OVRO observations of comet C/1999 S4 (LINEAR)

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### ABSTRACT

We present results from an observing campaign of the molecular content of the coma of comet C/1999 S4 (LINEAR) carried out jointly with the millimeter-arrays of the Berkeley-Illinois-Maryland Association (BIMA) and the Owens Valley Radio Observatory (OVRO). Using the BIMA array in autocorrelation ('single-dish') mode, we detected weak HCN  $J=1-0$  emission from comet C/1999 S4 (LINEAR) at  $14 \pm 4$  mK kms averaged over the  $143''$  beam. The three days over which emission was detected, 2000 July 21.9–24.2, immediately precede the reported full breakup of the nucleus of this comet. During this same period, we find an upper limit for HCN  $1-0$  of  $144 \text{ mJy beam}^{-1} \text{ km s}^{-1}$  (203 mK kms) in the  $9'' \times 12''$  synthesized beam of combined observations of BIMA and OVRO in cross-correlation ('imaging') mode. Together with reported values of HCN  $1-0$  emission in the  $28''$  IRAM 30-meter beam, our data probe the spatial distribution of the HCN emission from radii of 1300 to 19,000 km. Using literature results of HCN excitation in cometary comae, we find that the relative line

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fluxes in the  $12'' \times 9''$ ,  $28''$  and  $143''$  beams are consistent with expectations for a nuclear source of HCN and expansion of the volatile gases and evaporating icy grains following a Haser model.

*Subject headings:* radiative transfer – techniques: interferometric – comets: individual (C/1999 S4 LINEAR) – radio lines: solar system

## 1. Introduction

The volatile component of comets provides an excellent tracer of the chemical composition present during the formation of the Solar System (Mumma et al. 1993; Irvine & Bergin 2000; Irvine et al. 2000). Recent years have seen remarkable progress in this area with the appearances of two especially active comets, Hyakutake (C/1996 B2) and Hale-Bopp (C/1995 O1), that passed close to Earth. Observations at (sub) millimeter wavelengths of these spectacular comets, in addition to numerous less prominent but no less interesting comets over the last decade, indicate that the composition of comets is heterogeneous. This holds not only from comet to comet (e.g., A'Hearn et al. 1995), but also within the same comet as was shown with millimeter-interferometers that resolve the emission from the coma (e.g., Blake et al. 1999). The latter work illustrated the information that high spatial resolution observations can offer about cometary composition and evaporation processes.

In the Summer of 2000, the comet C/1999 S4 (comet LINEAR from here on) passed close to the Sun (0.765 AU) and at 0.374 AU to the Earth (Durig et al. 1999). Its orbit identified it as a dynamically young comet, and it showed a high level of activity at large heliocentric distance. As is not uncommon in dynamically young comets, the increase in its activity lagged significantly behind what was expected after its discovery, reaching only a medium level of activity at perihelion passage. It generated some excitement when, at closest approach to the Sun, the nucleus of the comet disintegrated and the residual material faded away in a few days, first reported by Kidger (2000). The comet's characteristics and its breakup are discussed in detail by (Bockelée-Morvan et al. 2001; Farnham et al. 2001; Lisse et al. 2001; Mäkinen et al. 2001; Mumma et al. 2001; Weaver et al. 2001; Bonev et al. 2002; Altenhoff et al. 2002; Tozzi & Licandro 2002; Schulz & Stüwe 2002; Cochran & Cochran 2001; Jockers et al. 2001; Rodríguez et al. 2002; Kidger 2002, and Magee-Sauer et al. 2001). Of most relevance to our discussion here are: the conclusion that the comet's nucleus may have broken up already several weeks before it fully disintegrated in late July (July 23–24); the molecular-line observations reported by Bockelée-Morvan et al. including the HCN-production rate; and the report by Magee-Sauer et al. (2001) of infrared HCN line observations.

This Paper describes the results of an observing campaign of the volatile composition of comet LINEAR, carried out jointly by the millimeter arrays at Hat Creek (operated by the Berkeley-Illinois-Maryland Association) and Owens Valley (operated by the California Institute of Technology), both in California (§2). We present a weak detection of HCN  $J=1-0$  in auto-correlation mode

and upper limits on the emission in cross-correlation mode, and place limits on the spatial extent of the emission (§3). Using models of HCN excitation in comets from Crovisier (1987) (see also Bockelée-Morvan et al. 1984) we investigate whether the spatial extent matches expectations (§4). Our conclusions are summarized in §5. The Appendix discusses the methods used to combine data from the two millimeter arrays.

## 2. Observations

Observations of comet LINEAR were obtained between 2000 July 17 and 2000 July 27 at the millimeter interferometers of the Berkeley-Illinois-Maryland Association (BIMA)<sup>1</sup> at Hat Creek, California (121° 28' 18".49 West, 40° 49' 2".5 North; altitude 1043 m), and the Caltech Millimeter Array at the Owens Valley Radio Observatory (OVRO)<sup>2</sup> near Big Pine, California (118° 16' 55".92 West, 37° 14' 2".04 North; altitude 1222 m). The six 10-meter antennas of the OVRO array were in ‘compact’ configuration, the ten 6-meter antennas of the BIMA array in ‘C’ configuration: both configurations result in comparable projected baselines: 2–20 k $\lambda$  at BIMA and 3–15 k $\lambda$  at OVRO. Since these were summer-time observations, when atmospheric stability can be poor for millimeter-interferometry, these compact configurations were adopted to avoid decorrelation by atmospheric phase fluctuations on long baselines.

OVRO observed the comet in crosscorrelation mode from rise in the early morning local time until set in the late afternoon local time. BIMA operated in crosscorrelation mode from rise until mid-morning local time. After mid-morning the phase stability at Hat Creek deteriorated significantly during the observing period, and the observing mode was switched to autocorrelation. This observing mode does not suffer from phase-decorrelation but has a large beam-dilution in the 143'' full-width at half-maximum (FWHM) primary beam of the six-meter antennas. The phase stability at Owens Valley, approximately 500 km south-east of Hat Creek and on the opposite side of the mountain range of the Sierra Nevada, was significantly better compared to Hat Creek, and successful crosscorrelation data were obtained throughout each day. In addition, the digital correlator at OVRO cannot operate in autocorrelation mode, which fortunately was not a limitation.

Receivers and digital correlators of both instruments were configured to record the lines of HCN  $J=1-0$  (88.6318470 GHz) and HCO<sup>+</sup>  $J=1-0$  (89.1885230 GHz). At OVRO each of the HCN and HCO<sup>+</sup> line frequencies were recorded in a single band of 64 channels with a resolution of 125 kHz (0.42 km s<sup>-1</sup>). At BIMA the lines were recorded in two bands per line, centered on the line frequency, with 128 and 256 channels each and with respective channels widths of 391 and

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98 kHz (1.3 and 0.33 km s<sup>-1</sup>). The autocorrelation mode of the BIMA correlator offers twice the number of channels at half the width for each of the bands, giving velocity resolutions of 0.65 and 0.17 km s<sup>-1</sup>.

Great care was taken to ensure that both arrays were using the same comet ephemeris at all times. On each of the observing days the most recent orbit solution from the Jet Propulsion Laboratory Solar System Dynamics Group<sup>3</sup> was used (Table 1). Changes of the predicted comet’s position between subsequent orbit solutions were less than a few arcseconds, smaller than the synthesized beam of  $\sim 12''$  and much smaller than the primary beam of  $143''$  for the BIMA autocorrelation observations. The small day-to-day uncertainty in the comet’s position therefore has no appreciable effect on the data. Because OVRO and BIMA use different coordinate systems (B1950.0 vs. J2000, respectively) and use, respectively, geocentric and topocentric representations for the coordinates of Solar-System objects, much effort was spent to ensure that both instruments actually pointed at the same position on the sky. Note that the parallax of the comet at a geocentric distance of 0.3 AU amounts to  $\sim 2''$  over the distance between the two arrays. The topocentric velocity of the comet was also Doppler tracked at both sites, keeping the cometocentric velocity centered at 0 km s<sup>-1</sup> in the spectral bands.

The complex gains of crosscorrelation data were calibrated separately for the BIMA and OVRO data, although the same flux and phase calibrators were used for both data. Depending on the position on the sky of the comet, the quasars 0716+714, 0923+392, and 1156+295 were used, with respective fluxes of 3.0, 1.2, and 1.2 Jy determined from contemporaneous measurements of planets at OVRO (Table 1). Using the same calibrator at each observatory ensures that the flux scaling of both data sets is identical. The OVRO data were calibrated using the ‘MMA’ package specific to Owens Valley (Scoville et al. 1993); the BIMA data were calibrated using the ‘MIRIAD’ software package (Sault et al. 1995). The OVRO data, originally stored in FITS format were subsequently converted to MIRIAD format, and both crosscorrelation data sets were added together after regridding to a common spectral resolution. The Appendix describes the data-combination procedure in more detail, including a test case validating the process. The combined data were then inverted and imaged using the MIRIAD software package.

The autocorrelation observations were corrected for atmospheric opacity and fluctuations by switching to an ‘off’ position at +20’ in Right Ascension every 30 seconds. The data were recorded on the main-beam brightness temperature scale in K; subsequent observations of the Galactic SiO maser R Leo in cross- and autocorrelation mode confirmed the flux scaling. A number of bad channels were flagged due to spikes, and second- to fourth-order polynomial baselines were subtracted from the autocorrelation spectra from the individual antennas. Finally the spectra from the individual antennas were added to yield a single autocorrelation spectrum for each of the spectral bands. All processing of the autocorrelation data was carried out within the MIRIAD software package.

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<sup>3</sup>Available at <http://ssd.jpl.nasa.gov/horizons.html>

### 3. Results

#### 3.1. Autocorrelation spectra

The autocorrelation spectra of HCN 1–0 (Fig. 1) and  $\text{HCO}^+$  1–0 were inspected using the average data of single, individual days and averaging over two, three, and six consecutive days. In HCN a weak signal was detected in two- and three-day averages between 2000 July 21.85 and 2000 July 24.16. These are the three days leading up to the disruption of the nucleus between 2000 July 24.9 and 25.9 (Kidger 2000) and coincide with a reported brightening at visible wavelengths, likely indicating a minor outburst event. The three-day average main-beam brightness temperature of HCN 1–0 is  $14 \pm 4 \text{ mK km s}^{-1}$  in the  $143''$  FWHM beam of the six-meter dishes. The emission is centered around a velocity of  $-1 \text{ km s}^{-1}$ , close to the expected velocity of the comet (nominally at zero), and has a width of  $\sim 2 \text{ km s}^{-1}$  FWHM. The latter value is uncertain because of the low signal-to-noise of the emission; the hyperfine components at  $+4$  and  $-7 \text{ km s}^{-1}$  were likewise below the noise level. The reported intensity is a total for all three components, obtained by multiplying the integrated intensity of the single detected component by the expected relative intensities in thermodynamic equilibrium:  $0.2 + 1.0 + 0.6 = 1.8$ . Although the overall signal-to-noise of the detection is low, the three-day averages leading up to (2000 July 21.85–24.16;  $14 \pm 4 \text{ mK km s}^{-1}$ ) and following (2000 July 24.95–27.17;  $< 10 \text{ mK km s}^{-1}$   $2\sigma$  upper limit) the comet’s final breakup around 2000 July 23–24 clearly show a dearth of the HCN 1–0 emission. The decrease in HCN  $J=3-2$  emission observed with the IRAM 30-meter telescope by Bockelée-Morvan et al. (2001) after 2000 July 23.8 supports our detected fall-off of HCN. No emission was detected in any of the  $\text{HCO}^+$  spectra, to a  $2\sigma$  upper limit of  $4 \text{ mK km s}^{-1}$  in three-day averages.

#### 3.2. Crosscorrelation images

The crosscorrelation data were inspected in a similar way as the autocorrelation data. The data were integrated over a single day, over multiple consecutive days, and over 4-hour time intervals to search for transient features. This was done for the data from the two arrays individually as well as for the combined data set. In all cases the emission was integrated over  $2 \text{ km s}^{-1}$  around a cometocentric velocity that was varied from  $-10$  to  $+10 \text{ km s}^{-1}$  to account for possible small tracking offsets in the comet’s velocity. The reported flux (limits) refer to the sum of the hyperfine components, added with the appropriate weights (1:5:3).

None of the images obtained in this way revealed any positive detection of HCN 1–0 or  $\text{HCO}^+$  1–0 emission. Figure 2 illustrates this for HCN 1–0 at an adopted cometocentric velocity of  $0 \text{ km s}^{-1}$  and including all OVRO and BIMA data. Using the three-day average over 2000 July 21.85–24.16, over which emission was detected in autocorrelation mode, an upper limit to the HCN 1–0 emission is found of  $144 \text{ mJy beam}^{-1} \text{ km s}^{-1}$  ( $2\sigma$ ) in the  $12'' \times 9''$  synthesized (naturally weighted) beam, corresponding to  $203 \text{ mK km s}^{-1}$ . The corresponding values for  $\text{HCO}^+$  1–0 are smaller, at

80 mJy beam<sup>-1</sup> km s<sup>-1</sup> and 117 mK km s<sup>-1</sup>, because HCO<sup>+</sup> has no hyper-fine splitting to sum over.

#### 4. Comparison to model calculations

The previous section showed that HCN 1–0 emission was detected, at low signal-to-noise, in autocorrelation mode but not in crosscorrelation mode. If the emission is unresolved ( $< 10''$ ), the beam-averaged value of  $14 \pm 4$  mK km s<sup>-1</sup> detected in the  $143''$  primary beam corresponds to an intensity of  $2.7 \pm 0.8$  K km s<sup>-1</sup> in the  $12'' \times 9''$  synthesized beam of the crosscorrelation data. Since we obtained an upper limit of 203 mK km s<sup>-1</sup>, it follows that the emission is significantly resolved out in the interferometer beam. Alternative explanations for reduced flux in the crosscorrelation data are decorrelation due to uncorrected phase variations and positional uncertainties of the comet larger than the synthesized beam. These processes are unlikely to affect the data: the measured decorrelation on the phase calibrator is typically less than 20%, and, as mentioned in §2, the position of the nucleus was known to within a few arcseconds from day to day, which is much smaller than the synthesized beam.

Extended emission is expected from molecules in the expanding coma. For a parent species like HCN, the density in the coma follows a Haser distribution,

$$n(r) = \frac{Q}{4\pi V_0 r^2} \exp\left(-\frac{r - r_n}{r_\lambda}\right), \quad (1)$$

where  $n$  is the number density of molecules,  $Q$  the production rate of that molecule by the nucleus,  $V_0$  is the expansion velocity in the coma,  $r_n$  is the size of the nucleus, and  $r_\lambda$  is characteristic photodissociation scale length. Bockelée-Morvan et al. (2001) quote  $Q(\text{HCN}) = 3 \times 10^{25}$  s<sup>-1</sup>; we adopt  $V_0 = 0.9$  km s<sup>-1</sup> and  $r_\lambda = V_0/\beta = 3.7 \times 10^4$  km, where  $\beta = 2.45 \times 10^{-5}$  s<sup>-1</sup> is the HCN photodissociation rate for the Solar spectrum at the time of the observations (Bockelée-Morvan 2001, private communication).

Is a distribution of HCN following eq. (1) consistent with our detection in the  $143''$  autocorrelation beam and our upper limit in the  $12'' \times 9''$  crosscorrelation beam? Or is an additional source of HCN required, distributed over  $> 10''$ , e.g., evaporating icy fragments that have broken off the nucleus? Fragments exceeding  $\sim 1$  m, non-uniformly illuminated by the Sun, are expected to undergo non-gravitational forces and accelerate away from the inner ( $\sim 10''$ ) coma (e.g., Desvoivres et al. 1999). Our data uniquely probe the spatial distribution of the HCN 1–0 emission. Bockelée-Morvan (2001, private communication) report HCN 1–0 emission in the  $28''$  IRAM 30-meter beam of  $93 \pm 13$  mK km s<sup>-1</sup>, averaged over 2000 July 21–23; we include this data point in our analysis. We do not consider the HCN 3–2 emission reported by Bockelée-Morvan et al. (2001) here, because we limit our discussion to the 1–0 transition.

The emission distribution of HCN 1–0 depends on the spatial distribution of the molecules following eq. (1) and on their excitation as function of position. The latter is given by the bal-

ance of (de-) excitation through collisions with water and electrons, excitation into vibrationally excited levels through absorption of Solar infrared photons, and de-excitation through spontaneous emission. Bockelée-Morvan et al. (1984) and Crovisier (1987) discuss the excitation of HCN in comets in detail. Instead of carrying out a full excitation calculation, we will use the HCN level populations modeled by Crovisier (1987): this paper presents three cases: water production rate  $Q(\text{H}_2\text{O})=2 \times 10^{29} \text{ s}^{-1}$  and heliocentric distance  $r_h = 1.0 \text{ AU}$  (model 1);  $Q(\text{H}_2\text{O})=2 \times 10^{28} \text{ s}^{-1}$  and  $r_h = 1.6 \text{ AU}$  (model 2); and  $Q(\text{H}_2\text{O})=1 \times 10^{30} \text{ s}^{-1}$  and  $r_h = 0.8 \text{ AU}$  (model 3). None of the parameters of these models are exactly identical to those for comet LINEAR ( $Q(\text{H}_2\text{O})=3 \times 10^{28} \text{ s}^{-1}$ , Bockelée-Morvan et al. 2001; and  $r_h = 0.76 \text{ AU}$ ), but we show that especially the *relative* line intensities in the  $12'' \times 9''$ ,  $28''$ , and  $143''$  beams are comparatively insensitive to the model details and are largely determined by the underlying column density distribution. Although the exact values of the of the relative line intensities in the three beams depend on the details of the excitation distribution in the coma that results from the balance between fluorescent and collisional processes, values within the same range are found in the extreme cases where either of these processes dominate. The models described above can therefore be used to test if our data agree with a Haser distribution for the column density.

Using a Haser distribution following eq. (1) with the parameters listed above, and the  $J=0$  and  $J=1$  HCN level populations of the three models from Crovisier (1987), we calculate the expected HCN 1–0 emission distributions at a distance of 0.374 AU from Earth. We convolve the emission distributions with Gaussians of  $143''$  and  $28''$  FWHM to get the main-beam intensities in these beams. For the crosscorrelation data, we construct synthetic visibilities based on the predicted emission distribution and the actual antenna positions of the data. This explicitly takes into account the spatial filtering of the interferometers and their insensitivity to large-scale structure, as well as the data combination procedure outlined in the Appendix. We process the synthetic visibilities in the same way as the data to obtain integrated-intensity images. We find that no more than 10–30% of the flux is resolved out by the interferometer.

Table 2 lists the observational results in the  $143''$ ,  $28''$ , and  $12'' \times 9''$  beams, alongside the predicted intensities in these beams for the three models. In terms of absolute intensities the observed values are within a factor of two of the model results, indicating a general agreement between the models and the observations. Table 2 also lists the observed and modeled intensities normalized to the value in the  $28''$  beam:  $F_{\text{auto}}/F_{28}$  and  $F_{\text{cross}}/F_{28}$ . For the three models,  $F_{\text{auto}}/F_{28}$  ranges from 0.21 to 0.28. Our observed value,  $0.15 \pm 0.05$ , is consistent with the models at the  $\sim 1\sigma$  level. Given the possibility of calibration differences between the IRAM 30-meter data and the BIMA data, we conclude that this indicates that the line emission over  $143''$  is consistent with the expectation of a Haser distribution when compared to a  $28''$  region. Bockelée-Morvan et al. (2001) estimate that only fragments with sizes  $\gtrsim 1 \text{ m}$  are large enough to survive outside the inner few thousand km of the coma. Our  $143''$  result shows that such fragments, if they exist, do not carry a significant contribution of the volatile material ( $\lesssim 50\%$ ). This is consistent with the results of Magee-Sauer et al. (2001), who report that spatially resolved infrared HCN line observations

are consistent with a nuclear source, and of Bockelée-Morvan et al. (2001), who conclude that the evolution of the HCN production rate suggests that HCN is mainly released by icy fragments of small size.

The ratios  $F_{\text{cross}}/F_{28}$  of the three models show a larger range, 1.57–3.65, compared to  $F_{\text{auto}}/F_{28}$ . This indicates that on these smaller scales, excitation differences exist depending on the heliocentric distance and evaporation rate (= density of collision partners). Our non-detection of emission in the crosscorrelation beam yields a strict upper limit  $F_{\text{cross}}/F_{28} < 2.5$ . This is consistent with the range predicted by the models, especially when also taking into account possible decorrelation of our crosscorrelation data up to 20% (see above) and, again, possible calibration differences between the IRAM 30-meter and OVRO+BIMA data. In other words, the non-detection in our combined crosscorrelation data is expected for a Haser-model HCN distribution. This indicates, that to obtain high-spatial resolution (sub) millimeter observations of moderately active comets such as LINEAR, higher sensitivity is required. The Combined Array for Research in Millimeter Astronomy (CARMA), planned to be operational in 2005, will already provide the increase in sensitivity required for a marginal detection of a comet like LINEAR; with the advent of the Atacama Large Millimeter Array (ALMA) such objects will become easily accesible. Needless to say, detailed excitation calculations along the lines of Bockelée-Morvan et al. (1984), Crovisier (1987), and Biver et al. (1999) will be needed to interpret these future observations.

## 5. Conclusions

1. We detected HCN 1-0 emission in autocorrelation mode with the BIMA array at  $14 \pm 4$  mK km s<sup>-1</sup> in the 143'' beam, and find a  $2\sigma$  upper limit of 144 mJy beam<sup>-1</sup> km s<sup>-1</sup> (203 mK km s<sup>-1</sup>) in the 12'' × 9'' synthesized beam of combined crosscorrelation data from OVRO and BIMA. Only upper limits are obtained for HCO<sup>+</sup> of < 4 mK km s<sup>-1</sup> in autocorrelation and < 80 mJy beam<sup>-1</sup> km s<sup>-1</sup> in crosscorrelation. HCN emission was detected in the three days preceding (2000 July 21.9–24.2) the breakup of the comet's nucleus around 2000 July 23–24; no emission in HCN or HCO<sup>+</sup> was detected after the breakup.
2. Compared to models of HCN excitation in cometary comae, we find that the relative and absolute integrated intensities in 143'', 28'', and 12'' × 9'' are consistent with expectations. This indicates a nuclear source for HCN, and no significant contribution from a spread-out (> 1300 km) distribution of evaporating icy fragments. Fragments exceeding 1 m in size are thought to be able to resist evaporation over such distances, and our data limit their contribution to the volatiles to  $\lesssim 50\%$  during the days of 2000 July 21–23 immediately preceding the full disruption of comet LINEAR.
3. We showed that data obtained at the BIMA and OVRO interferometers can be successfully combined, creating a 'virtual' CARMA array that increases the signal-to-noise by approximately a factor of  $\sqrt{2}$  due to the increased collecting array. While for these observations



the array configurations were chosen to provide overlapping, short  $uv$  spacings because of atmospheric phase-stability concerns, virtual ‘CARMA’ observations have the possibility to offer complementary array configurations increasing imaging capabilities. These, however, fall short of what a truly combined array with baselines between the individual BIMA and OVRO antennas can offer: a combined CARMA array is foreseen to be operational by 2005. With the arrival of the Atacama Large Millimeter Array, spatially resolved comet observations at (sub) millimeter wavelengths will become increasingly accesible.

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### A. Combining data from different interferometer arrays

The results presented in this Paper include images that are constructed based on data obtained simultaneously at two different millimeter interferometers, BIMA and OVRO. The process of combining such data contains a number of steps where special care needs to be taken to ensure that the resulting images correctly represent the data obtained at the individual arrays. This Appendix describes in detail the data combination process using the MIRIAD software package, and presents simulations that validate the procedure (Fig. 3).

As a first step, even before the data are taken, one needs to ensure that both arrays are using the same source position as phase tracking system, and that the receivers and correlators are configured to record the same line frequency using the same systemic velocity of the object at all times. Especially for a moving target like a comet, these requirements are not trivial when dealing with two separate observatories.

After the data have been taken, the time variation of the complex gains (amplitude and phase) has to be calibrated for the two data sets separately since there is no correlation between the gain variations at the two sites. It is not necessary that both arrays use the same objects, usually quasars, as phase and amplitude calibrators. For the absolute calibration of the flux, however, it can be advantageous to include the same secondary calibrator (a bright quasar), the flux of which is derived from contemporaneous observations of a primary calibrator (planet), preferably with both

arrays. This ensures that there are no offsets in flux calibration between the arrays. Especially if the arrays have a different response to spatial scales, such flux offset would affect the resulting combined image.

Before the calibrated data sets can be combined, they need to be converted to a common data format (if not already the same). The OVRO data are usually exported in FITS format, while BIMA data are written in a format particular to the MIRIAD software package. The MIRIAD task ‘FITS’ can convert FITS-files to MIRIAD-files. To successfully convert OVRO data using this task, use ‘`options=varwt`’ so that the weights of the visibility data written by MMA (the OVRO data calibration package) are properly interpreted as the reciprocal of the noise variance. Further processing in the MIRIAD package does not require any additional conversions: many MIRIAD tasks accept input data sets that, for example, have different spectral gridding as do our OVRO and BIMA data.

Figure 3 illustrates the data combination procedure. To test our combination procedure, this figure also shows the results of combining data with a relative spatial offset: this results in an image with two ‘sources’, each with approximately 50% of the flux because in our observations the sensitivity of the OVRO and BIMA data was about equal.

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Fig. 1.— Autocorrelation spectra obtained in the  $143''$  BIMA primary beam. The top row shows single-day averages; the subsequent rows show the two-day, three-day, and six-day averages. Emission of HCN 1–0 is detected at  $\sim 4\sigma$  averaged over 2001 July 21.65–24.16, indicated by the grey bar in the corresponding panel.

Fig. 2.— Crosscorrelation images of the combined BIMA and OVRO data centered on a velocity of  $0 \text{ km s}^{-1}$ , at  $2\sigma$  contours ( $1\sigma=42 \text{ mJy beam}^{-1} \text{ km s}^{-1}$  for the single-day averages). The top row shows the single-day averages; the subsequent rows show the averages over two, three, and six days. The synthesized beam is indicated in the lower left corner of each panel.

Fig. 3.— The top two panels show the model image, as observed by OVRO (left) and BIMA (right). After combination, the ‘virtual CARMA’ result is shown at the lower left. By introducing an offset in the source position of  $+10''$  for the BIMA model and  $-10''$  for the OVRO model, the resulting image in the combined data (lower right) is split. This latter panel shows that the data combination procedure in fact preserves the data from both arrays.

Table 1. Observations

OVRO cross corr. UT 2000 July	BIMA cross corr. UT 2000 July	BIMA auto corr. UT 2000 July	JPL orbit solution#	Gain calibrator <sup>a</sup>	Calibrator's flux (Jy)
18.47–19.15	...	...	62	0716+714	3.0
19.47–19.88	...	...	68	0716+714	3.0
19.95–19.99	...	...	68	0716+714	3.0
20.50–20.93	20.45–20.73	...	77	0716+714	3.0
21.55–22.13	21.69–21.82	21.85–22.13	77	0923+392	4.5
22.59–23.14	22.63–22.92	22.95–23.12	78	0923+392	4.5
23.64–24.08	23.65–23.92	23.95–24.16	82	0923+392	4.5
24.70–25.12	24.67–24.92	24.95–25.12	82	0923+392	4.5
25.72–26.16	25.69–25.94	25.99–26.30	83	0923+391	4.5
26.76–27.19	26.70–26.89	26.92–27.17	87	1156+295	1.2

<sup>a</sup>Source designation following B1950-based OVRO calibrator list,  
<http://www.ovro.caltech.edu/mm/preobs/guide.html>

Table 2. HCN 1–0 Intensities: Observations and Models

	BIMA autocorr. 143''	IRAM 30-m 28''	BIMA+OVRO crosscorr. 12'' × 9''
Integrated intensities, in mK km s <sup>−1</sup>			
Observed	14 ± 4	93 ± 13 <sup>a</sup>	< 203
Model 1 <sup>b</sup>	12.4	44.4	110.0
Model 2 <sup>c</sup>	22.7	110.0	172.8
Model 3 <sup>d</sup>	5.96	24.6	89.9
Integrated intensities, normalized to 28'' beam			
Observed	0.15 ± 0.05	≡ 1	< 2.5
Model 1	0.28	≡ 1	2.15
Model 2	0.21	≡ 1	1.57
Model 3	0.24	≡ 1	3.65

<sup>a</sup>Bockelée-Morvan 2000, priv. comm.

<sup>b</sup> $Q(\text{H}_2\text{O})=2 \times 10^{29} \text{ s}^{-1}$ ,  $r_h = 1.0 \text{ AU}$ , from Crovisier (1987), their Fig. 2c.

<sup>c</sup> $Q(\text{H}_2\text{O})=2 \times 10^{28} \text{ s}^{-1}$ ,  $r_h = 1.6 \text{ AU}$ , from Crovisier (1987), their Fig. 3c.

<sup>d</sup> $Q(\text{H}_2\text{O})=1 \times 10^{30} \text{ s}^{-1}$ ,  $r_h = 0.8 \text{ AU}$ , from Crovisier (1987), their Fig. 4c.

Note. — Observed line intensities and upper limit are averages over 2000 July 21–23.  $Q(\text{HCN}) = 3 \times 10^{25} \text{ s}^{-1}$  has been adopted for the calculations, following Bockelée-Morvan et al. (2001).

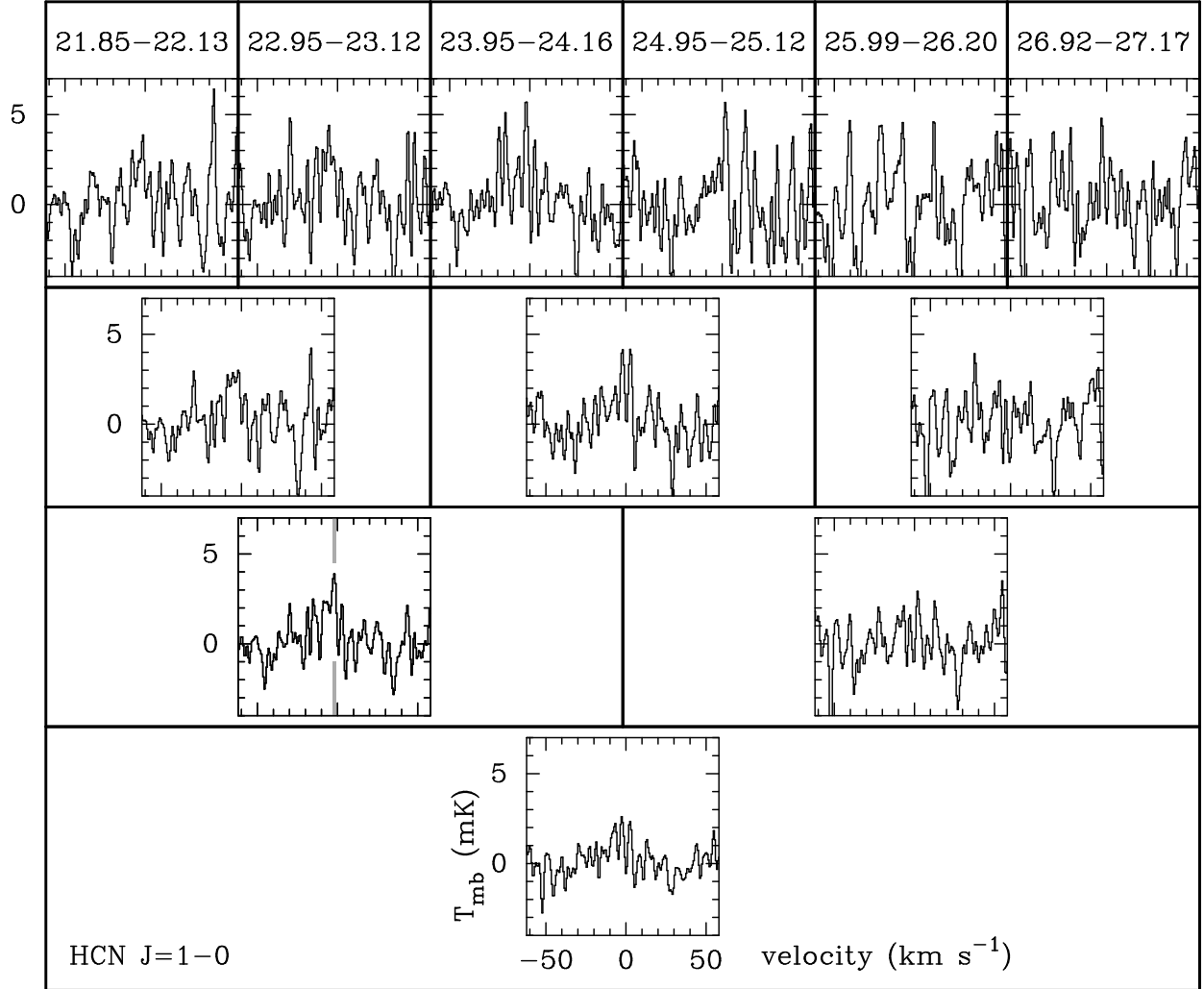


Fig. 1.—



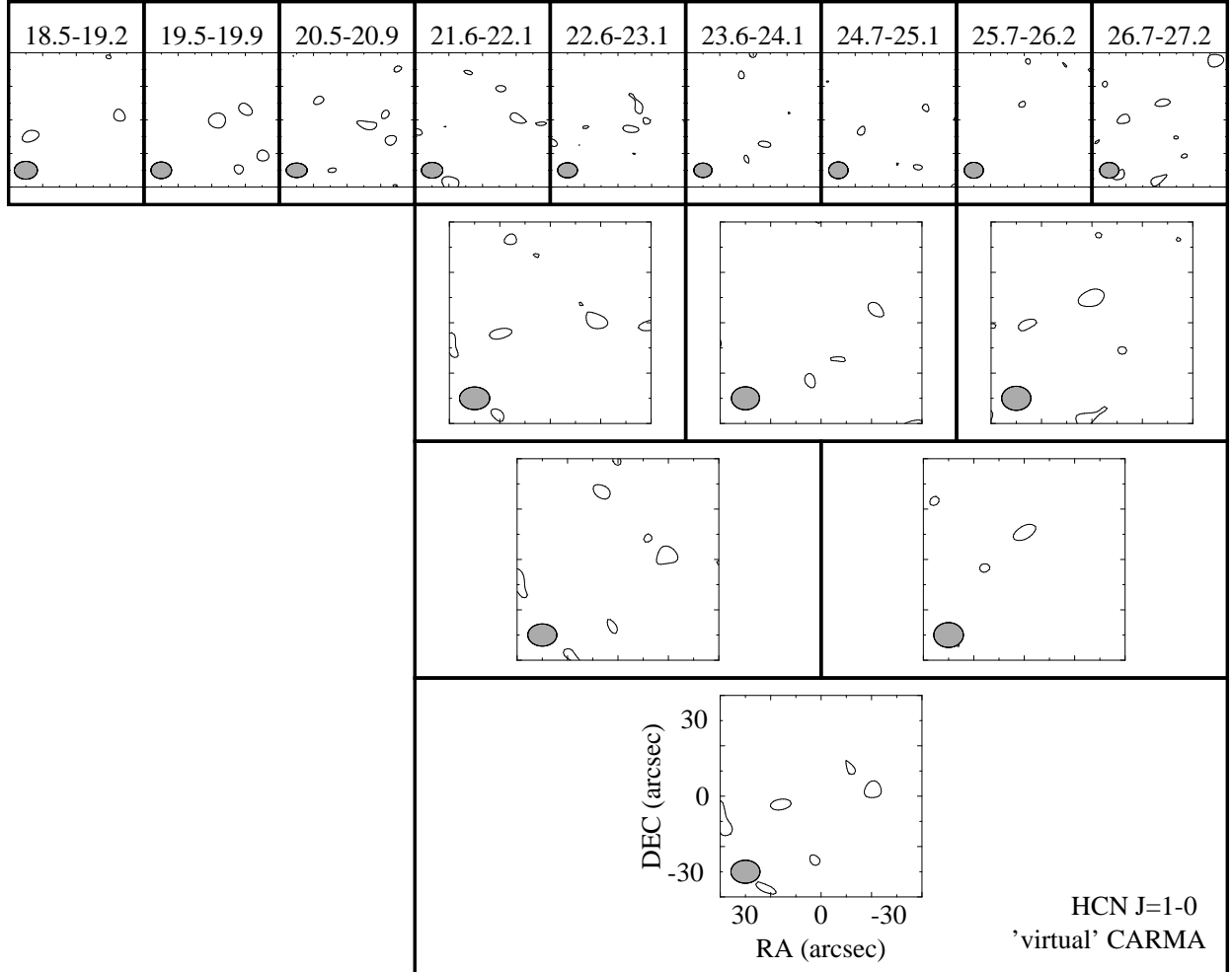


Fig. 2.—

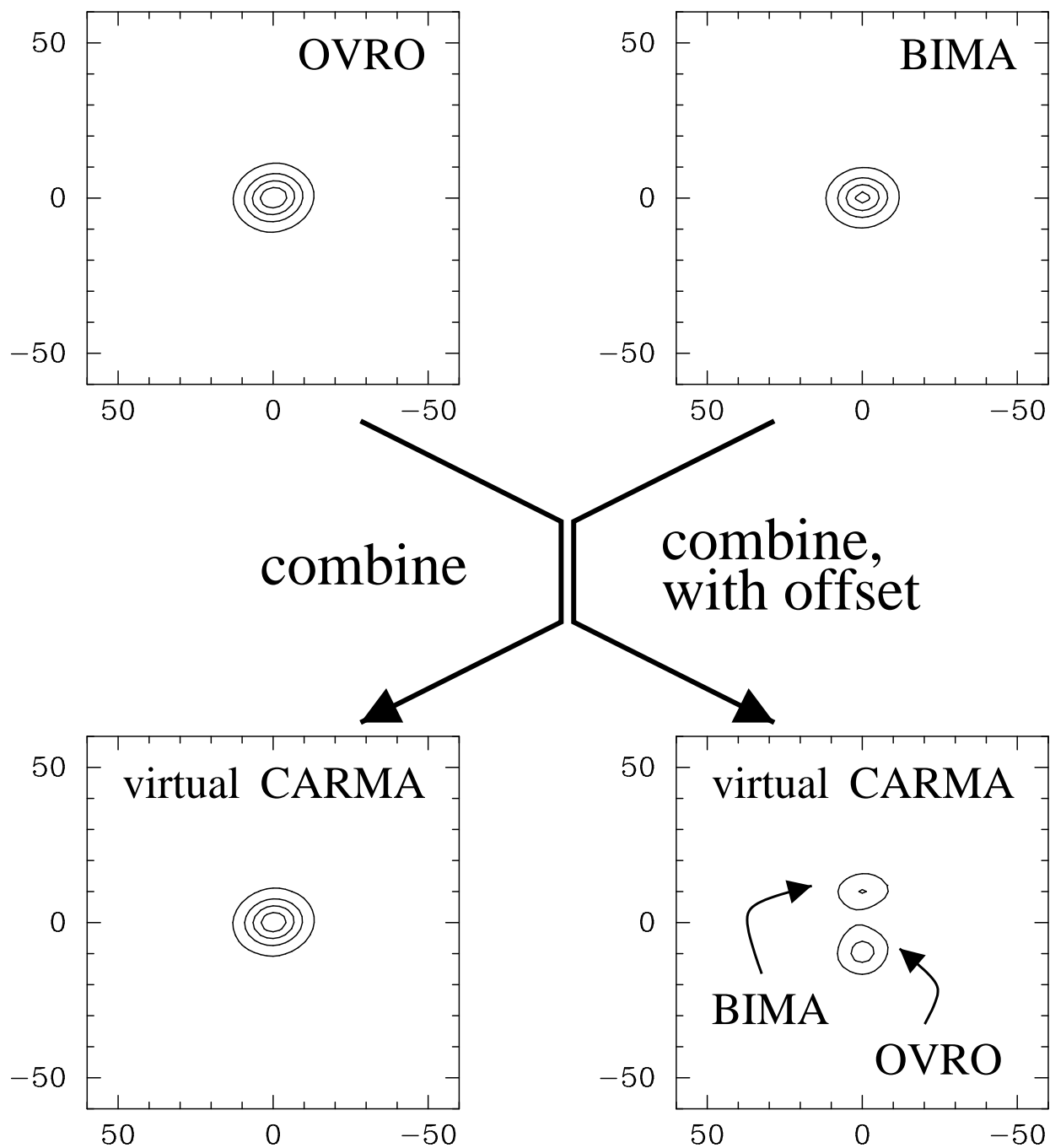


Fig. 3.—